

## SAGE-VAR: AN INFRARED SURVEY OF VARIABILITY IN THE MAGELLANIC CLOUDS

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Received 2015 February 21; accepted 2015 April 27; published 2015 June 24

## ABSTRACT

We present the first results from the Surveying the Agents of Galaxy Evolution (SAGE)-Var program, a follow up to the *Spitzer* legacy program SAGE (Meixner et al.). We obtained four epochs of photometry at 3.6 and 4.5  $\mu\text{m}$  covering the bar of the LMC and the central region of the SMC in order to probe the variability of extremely red sources missed by variability surveys conducted at shorter wavelengths, and to provide additional epochs of observation for known variables. Our six total epochs of observations allow us to probe infrared (IR) variability on 15 different timescales ranging from  $\sim 20$  days to  $\sim 5$  yr. Out of a full catalog of 1 717 554 (LMC) and 457 760 (SMC) objects, we find 10 (LMC) and 6 (SMC) large amplitude Asymptotic Giant Branch (AGB) variables without optically measured variability owing to circumstellar dust obscuration. The catalog also contains multiple observations of known AGB variables, type I and II Cepheids, eclipsing variables, R CrB stars, and young stellar objects, which will be discussed in following papers. Here we present IR Period–Luminosity (PL) relations for classical Cepheids in the Magellanic Clouds, as well as improved PL relationships for AGB stars pulsating in the fundamental mode using mean magnitudes constructed from six epochs of observations.

**Key words:** stars: AGB and post-AGB – stars: mass-loss – stars: variables: general – stars: variables: Cepheids – Magellanic Clouds

## 1. INTRODUCTION

The study of stellar variability has had a wide-ranging impact on astronomy and cosmology. Recent variable star surveys such as the MASSive Compact Halo Object search (MACHO; Alcock et al. 1997) and the Optical Gravitational Lensing Experiment (OGLE; Udalski et al. 1997) have generated catalogs of tens of thousands of variable stars of numerous classes in the Magellanic Clouds. However, like most ground-based variability surveys, both of these were performed at visible wavelengths, and therefore miss the reddest variable sources, such as dust-enshrouded, highly evolved Asymptotic Giant Branch (AGB) stars, which are nearly invisible except in the infrared (IR). AGB stars are unstable, and exhibit variability on timescales of hundreds of days (Vassiliadis & Wood 1993). Due to this, they are also classified as long-period variables (LPVs). There are several surveys that monitored LPVs at near-IR wavelengths in the Magellanic Clouds (*J*, *H*, and *K*-bands; Ita et al. 2002; Whitelock et al. 2003), but even these can still miss the dustiest sources. The pulsation properties of the heavily enshrouded (and most evolved) AGB stars are therefore not well understood. This problem can be addressed with IR monitoring at wavelengths longer than 3  $\mu\text{m}$ , but there are very few examples of mid-IR monitoring surveys in any galaxy. Le Bertre (1992, 1993) obtained light curves at 1–20  $\mu\text{m}$  of  $\sim 60$  O- and C-rich AGB stars in our own galaxy and found that while pulsation amplitudes generally decrease into the mid-IR, circumstellar dust can cause amplitudes to increase again at

$\lambda > 3 \mu\text{m}$ . In M33, McQuinn et al. (2007) obtained five epochs of imaging at 3.6, 4.5, 5.8, and 8  $\mu\text{m}$  and found that the pulsation amplitude tends to increase with color, but that this relationship may break down at the longest wavelengths. In the Magellanic Clouds and other nearby dwarf galaxies, only two to three epochs of imaging are available at  $\lambda > 3 \mu\text{m}$  (Vijh et al. 2009; Boyer et al. 2015a, 2015b; Polsdofer et al. 2015). While these surveys can detect flux changes in a large fraction of the dustiest stars, they cannot place stringent limits on the pulsation periods or amplitudes in the IR. In order to explore the variability of these reddest sources, which dominate the mass return to the interstellar medium (ISM) from evolved stars (Riebel et al. 2012), we used the *Spitzer Space Telescope* to survey the bar regions of the LMC and SMC at 3.6 and 4.5  $\mu\text{m}$ . SAGE-Var represents the first large-scale variability survey at such red wavelengths. While our original focus was the reddest AGB stars, we have detected over 2700 IR variables in the Clouds of many classes.

This paper is organized as follows. In Section 2 we detail the observational strategy (2.1), data reduction (2.2), catalog construction (2.3), and source classification scheme (2.4) for the SAGE-Var program. In Section 3, we discuss our results: the identification of LPV candidates without prior observed variability (3.1), an investigation of the connection between variability amplitude and dust-production rate (DPR) from evolved stars (3.2), and new measurements of the IR Period–Luminosity (PL) relationship for both LPVs and classical Cepheids in the Magellanic Clouds (3.3). Our conclusions are presented in Section 4.

**Table 1**  
Observation Dates

Epoch	Date	Julian Date
LMC:		
Epoch 1:	2005 Jul 20	2453572
Epoch 2:	2005 Oct 28	2453672
Epoch 3:	2010 Aug 17	2455426
Epoch 4:	2010 Sep 10	2455450
Epoch 5:	2010 Dec 25	2455556
Epoch 6:	2011 Apr 27	2455679
SMC:		
Epoch 1:	2008 Jun 15	2454633
Epoch 2:	2008 Sep 19	2454729
Epoch 3:	2010 Aug 17	2455426
Epoch 4:	2010 Sep 12	2455452
Epoch 5:	2010 Dec 24	2455555
Epoch 6:	2011 Jun 16	2455729

**Note.** Dates of the six epochs of observations of the SAGE-Var project, including the two epochs of the original SAGE-LMC and SAGE-SMC programs (Epochs 1 and 2). The dates listed for the original two epochs of observations are the approximate midpoints of the  $\sim 5$  days observation periods, while later epochs were executed during a single 24 hr period.

## 2. THE DATA

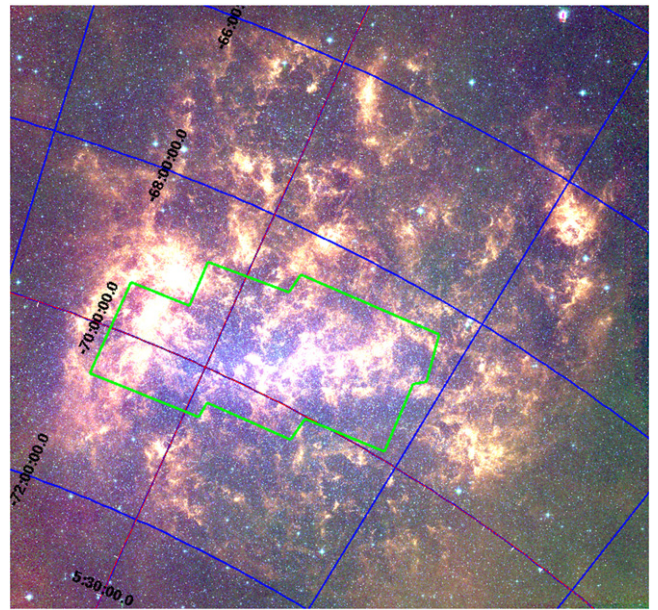
### 2.1. Observations

The four epochs of SAGE-Var observations (*Spitzer* PID 70020) were taken over a 10 month period, between August 2010 and June 2011 (see Table 1). Each epoch is a  $3.7 \times 1.5$  (LMC, Figure 1) or  $1.7 \times 1.7$  (SMC, Figure 2) mapping of the bar region of the galaxy, using both the [3.6] and [4.5] bands of the *Spitzer* warm mission. All frames were taken as 12 s exposures using the IRAC High Dynamic Range mode, which also produces short 0.6 s exposures in order to mitigate the effects of saturation for the brightest sources. This exposure mode was chosen to provide uniformity with the previous two epochs of the SAGE-LMC (Meixner et al. 2006) and SAGE-SMC (Gordon et al. 2011) surveys. The scheduling of our observations was determined using a Madore & Freedman (2005) power-law cadence with an index of 0.99, taking into account the original two SAGE epochs. Including these two original epochs, our six total epochs of observations allow us to probe variability on 15 different timescales (see Section 2.3).

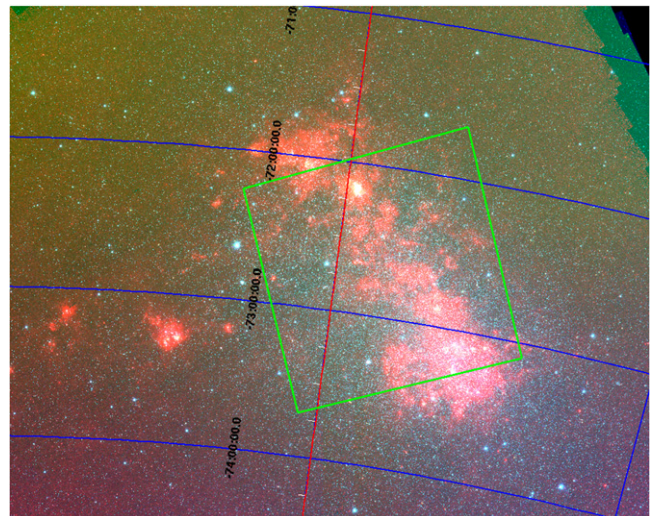
### 2.2. Data Reduction

The SAGE-Var data were processed using a pipeline developed at the University of Wisconsin for the GLIMPSE survey (Benjamin et al. 2003), which was also used for processing the original SAGE-LMC and SAGE-SMC observations. The Wisconsin pipeline corrects for numerous observational artifacts such as stray light, column pull-down, banding, and bad pixels. These observational artifacts are discussed by Hora et al. (2004) and the Wisconsin pipeline is described in Section 3.1.2 of Meixner et al. (2006). The pipeline is based on DAOPHOT (Stetson 1987) and is modified to account for the variable diffuse background in IRAC images. Photometry is performed on the Basic Calibrated Data (BCD) frames, and the FWHM of the point-spread function is  $\sim 1''.7$  in both bands.

In addition, the individual frames are mosaicked, and point sources identified. During catalog construction (Section 2.3), the point source lists are also position matched to 2MASS *JH*



**Figure 1.** Three-color image of the LMC (red:  $8.0 \mu\text{m}$ , green:  $5.8 \mu\text{m}$ , blue:  $3.6 \mu\text{m}$ ) showing the footprint of the SAGE-Var observations in green, focused on the stellar bar region of the galaxy.

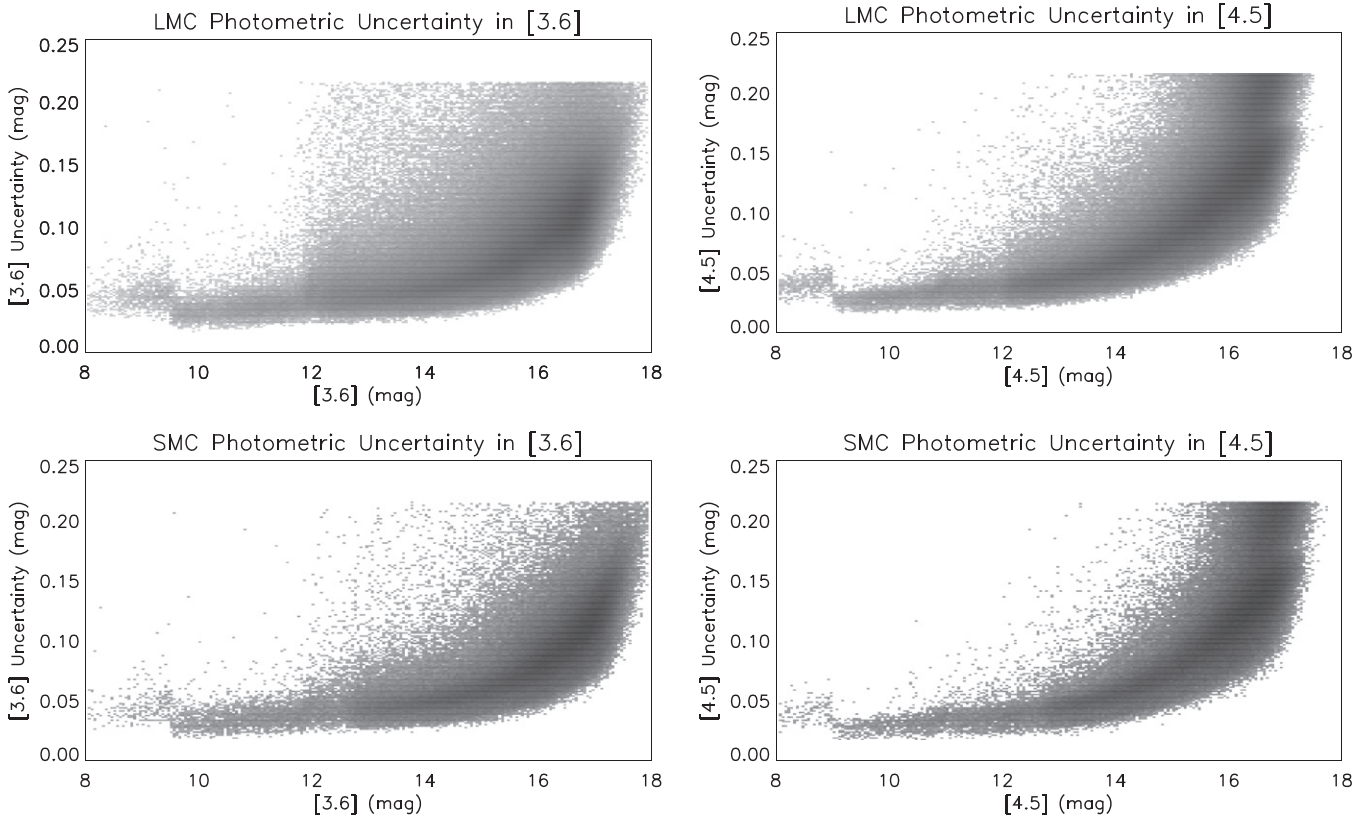


**Figure 2.** Three-color image of the SMC (red:  $24 \mu\text{m}$ , green:  $4.5 \mu\text{m}$ , blue:  $3.6 \mu\text{m}$ ) showing the footprint of the SAGE-Var observations in green, focused on the main stellar locus of the galaxy.

$K_s$  photometry. The 2MASS bands have been dereddened to account for interstellar extinction. The reddening coefficients used for the LMC can be found in Table 1 of Riebel et al. (2012). The photometry for the more distant SMC was dereddened using a value of  $E_{B-V} = 0.04$  mag (Harris & Zaritsky 2004; Schlegel et al. 1998) and the prescription of Glass et al. (1999, p 109). Our fluxes were transformed into magnitudes using zero points of 6.12 mag (280.9 Jy) for [3.6] and 5.63 mag (179.7 Jy) for [4.5] (Reach et al. 2005).

The pipeline provides  $1\sigma$  photometric uncertainty in each band, plotted as a function of source magnitude in Figure 3. Throughout this paper, we take these  $1\sigma$  errors as the uncertainties in our photometry. The distributions in Figure 3 show a hard cutoff at 0.2 mag due to the requirement that sources in the original SAGE survey have a signal-to-noise





**Figure 3.** Each plot shows the  $1\sigma$  uncertainty produced from the Wisconsin pipeline as a function of source magnitude for the entire SAGE-Var data set as a Hess diagram. The top row shows the data from the LMC ([3.6] on the left, [4.5] on the right), while the bottom row shows the SMC data.

ratio ( $S/N > 5$ ), which effectively results in the top cutoff. Sources brighter than magnitude  $\sim 9.0$  are observed only using the 0.6 s exposure mode, which results in a larger uncertainty than the 12 s exposure time. Between magnitudes 9 and 12, the 0.6 s and 12 s exposures are averaged, and the photometry of sources dimmer than magnitude  $\sim 12$  is derived only from the 12 s exposures. These three regimes of source brightness result in the breaks visible at  $\sim 9$  and  $\sim 12$  mag. See Section 3.1.3 of Meixner et al. (2006) for a full discussion.

### 2.3. Catalog Construction

The initial data product was a “full” point source list, without cosmic ray screening. These artifacts were removed from the catalog by matching the full lists to the SAGE mosaic photometry archive<sup>11</sup> using a  $2''$  matching radius. The mosaic archives of the original SAGE surveys were constructed by co-adding and re-reducing the photometry from the original two SAGE Epochs. The archives are deeper, more complete, and of higher  $S/N$  than the SAGE-Var data due to the greater exposure time of the mosaic photometry (up to  $\sim 50$  s per pixel compared to 12 s) and the fact that the original surveys were performed during the cold phase of the *Spitzer* mission whereas the SAGE-Var observations were taken after the liquid helium cryogen on board *Spitzer* was exhausted. The higher operating temperature of the *Spitzer* instruments restricts our observations to only the shortest IRAC wavelengths, and lends additional thermal noise to our images. Due to their greater depth and quality, the SAGE mosaic archives serve as a “truth field” for

the SAGE-Var full source lists, providing a more thorough list of actual astronomical sources while screening out instrumental artifacts by position matching. After the SAGE-Var source list was matched to the mosaic photometry, the original SAGE Epoch 1 and 2 archive data were matched individually as well, producing a final source list of 1 717 554 (LMC) and 457 760 (SMC) unique sources with up to six individual 3.6 and  $4.5\ \mu\text{m}$  observations, as well as mosaic photometry available in the longer wavelength IRAC bands. Table 2 lists the number of unique sources detected in each band of each epoch of SAGE-Var observations. Also listed is the number of objects detected in only one epoch, in two epochs, in three epochs, etc. A greater number of detections is preferable. The entire SAGE-Var data set, including all six epochs of observations in both bands, is available as an online table, hosted at the NASA/IPAC Infrared Science Archive.<sup>12</sup> Tables 3 and 4 provide a guide to the format of the online catalogs.

### 2.4. Source Classification

Variables in the SAGE-Var data set were identified using the variability criteria of Vihj et al. (2009). We calculate variability indices

$$V_{ij}^b = \frac{f_i^b - f_j^b}{\sqrt{(\sigma_i^b)^2 + (\sigma_j^b)^2}}$$

for every star, for every possible combination of epochs  $i$  and  $j$ , for the flux  $f$  in each SAGE-Var band  $b$ , ([3.6] and [4.5]) with

<sup>11</sup> [http://irsa.ipac.caltech.edu/data/SPITZER/SAGE/doc/SAGEDataProductsDescription\\_Sep09.pdf](http://irsa.ipac.caltech.edu/data/SPITZER/SAGE/doc/SAGEDataProductsDescription_Sep09.pdf)

<sup>12</sup> <http://irsa.ipac.caltech.edu/Missions/spitzer.html>

**Table 2**  
Object Count

Epoch	LMC		SMC	
	3.6 $\mu\text{m}$	4.5 $\mu\text{m}$	3.6 $\mu\text{m}$	4.5 $\mu\text{m}$
Mosaic Photometry	1,712,135	1,712,994	454,951	455,587
SAGE Epoch 1	1,185,774	1,083,380	297,732	269,178
SAGE Epoch 2	1,163,888	1,068,438	290,055	262,408
SAGE-Var 1	1,177,715	824,928	289,359	185,392
SAGE-Var 2	1,242,129	852,734	264,921	176,911
SAGE-Var 3	1,244,543	855,533	287,168	173,756
SAGE-Var 4	1,244,791	845,370	280,656	172,379
1 Detection	116,285	240,980	56,602	78,408
2 Detections	197,701	240,983	75,528	61,853
3 Detections	270,324	273,872	79,241	65,429
4 Detections	253,951	183,972	54,222	45,670
5 Detections	271,603	180,249	64,538	43,291
6 Detections	593,727	391,448	120,822	73,748
Total Sources	1,717,554 <sup>a</sup>		457,760 <sup>a</sup>	

**Notes.** Summary of the number of unique objects detected in the SAGE-Var survey of the SMC and LMC.

<sup>a</sup> These numbers represent the total number of unique sources detected in each galaxy, with an absolute minimum of one detection in one band.

**Table 3**  
Full Object Catalog Contents

Column	Name	Description	Null
1	desig	Source IRAC Designation	...
2	ra	Right Ascension, J2000 (deg)	...
3	dec	Declination, J2000 (deg)	...
4	e1_36	Epoch 1 flux in [3.6] band (Jy)	−99
5	e1_36_u	Uncertainty in Epoch 1 [3.6] flux (Jy)	−99
6–15	eN_36	Epoch 2–6 flux/uncertainty in [3.6] band <sup>a</sup> (Jy)	−99
16–27	eN_45	Epoch 1–6 flux/uncertainty in [4.5] band <sup>b</sup> (Jy)	−99
28	mean_36	Mean [3.6] flux (Jy)	−99
29	mean_36_u	rms Uncertainty in mean [3.6] flux (Jy)	−99
30	mean_45	Mean [4.5] flux (Jy)	−99
31	mean_45_u	rms Uncertainty in mean [4.5] flux (Jy)	−99
32–46	var_36_N	Variability index at [3.6] for interval <sup>c</sup> $N$	NaN
47–61	var_45_N	Variability index at [4.5] for interval <sup>c</sup> $N$	NaN
62	ogle_id	ID From the OGLE-III Catalog of Variable Stars	...
63	ogle_class	Classification from the OGLE-III CVS	...
64	ogle_per	Variability period from the OGLE-III CVS (days)	−99
65	macho_id	ID from the MACHO survey	...
66	macho_per	Variability period from the MACHO survey (days)	−99
67	grams_class	Classification of best-fitting GRAMS model (C or O)	...
68	yso_class	“Y” if a source is classified as a YSO candidate	...

**Notes.** The full SAGE-Var catalog of 1 717 554 (457 760) sources in the LMC (SMC) is available from IRSA. This table is provided as a guide to the online catalog’s structure and content.

<sup>a</sup> The epoch 2 through 6 [3.6] photometry follows the same format as columns 4 and 5.

<sup>b</sup> The [4.5] photometry follows the same format as the [3.6] photometry in columns 4–15.

<sup>c</sup> Defined in Section 2.4 and Table 5.

**Table 4**  
Variable Object Catalog Contents

Column	Name	Description	Null
1	desig	Source IRAC Designation	...
2	ra	Right Ascension, J2000 (deg)	...
3	dec	Declination, J2000 (deg)	...
4	e1_36	Epoch 1 flux in [3.6] band (Jy)	−99
5	e1_36_u	Uncertainty in Epoch 1 [3.6] flux (Jy)	−99
6–15	eN_36	Epoch 2–6 flux/uncertainty in [3.6] band <sup>a</sup> (Jy)	−99
16–27	eN_45	Epoch 1–6 flux/uncertainty in [4.5] band <sup>b</sup> (Jy)	−99
28	mean_36	Mean [3.6] flux (Jy)	−99
29	mean_36_u	rms Uncertainty in mean [3.6] flux (Jy)	−99
30	mean_45	Mean [4.5] flux (Jy)	−99
31	mean_45_u	rms Uncertainty in mean [4.5] flux (Jy)	−99
32–46	var_36_N	Variability index at [3.6] for interval <sup>c</sup> $N$	NaN
47–61	var_45_N	Variability index at [4.5] for interval <sup>c</sup> $N$	NaN
62	ogle_id	ID From the OGLE-III Catalog of Variable Stars	...
63	ogle_class	Classification from the OGLE-III CVS	...
64	ogle_per	Variability period from the OGLE-III CVS (days)	−99
65	macho_id	ID from the MACHO survey	...
66	macho_per	Variability period from the MACHO survey (days)	−99
67	grams_class	Classification of best-fitting GRAMS model (C or O)	...
68	yso_class	“Y” if a source is classified as a YSO candidate	...
69	amp_36	SAGE-Var observed [3.6] Amplitude	...
70	amp_45	SAGE-Var observed [4.5] Amplitude	...

**Notes.** This table extracts just those sources flagged as variable by the criteria of Section 2.4. It follows essentially the same format as Table 3 with the addition of the observed amplitudes of the variable objects. These are simply the difference between the brightest and dimmest magnitudes observed for the source, and represent a lower limit on the source’s full variability.

<sup>a</sup> The epoch 2 through 6 [3.6] photometry follows the same format as columns 4 and 5.

<sup>b</sup> The [4.5] photometry follows the same format as the [3.6] photometry in columns 4–15.

<sup>c</sup> Defined in Section 2.4 and Table 5.

photometric uncertainty  $\sigma$ . The photometric uncertainties are taken directly from the Wisconsin pipeline (Section 2.2). The variability index is thus the number of standard errors by which two epochs differ in brightness. Vihj et al. (2009) had only the first two epochs of the original SAGE survey, but five bands of photometry. In this study, our situation is reversed, in that we have only two bands of photometry, but  $\binom{6}{2} = 15$  possible epochal differences. Each epochal difference probes variability on a different timescale (Table 5), and we term each epochal difference an *interval*. In order for a source to be classified as a variable in a given interval, we require it to exhibit  $3\sigma$  flux variation in the same direction (that is, brightening or dimming) in both the [3.6] and [4.5] bands ( $|V_{ij}^b| \geq 3$  and the same sign in both bands).

The SAGE-Var sample has a total of 819 sources in common with the sample of Vihj et al. (2009). Of these, we independently identify 752 (92%) as variables using our own

**Table 5**  
Interval Timescales

Epochs	LMC Interval (days)	SMC Interval (days)	Label
Epoch 1–Epoch 2	100	96	1
Epoch 1–Epoch 3	1854	793	2
Epoch 1–Epoch 4	1878	819	3
Epoch 1–Epoch 5	1984	922	4
Epoch 1–Epoch 6	2107	1096	5
Epoch 2–Epoch 3	1754	697	6
Epoch 2–Epoch 4	1778	723	7
Epoch 2–Epoch 5	1884	826	8
Epoch 2–Epoch 6	2007	1000	9
Epoch 3–Epoch 4	24	26	10
Epoch 3–Epoch 5	130	129	11
Epoch 3–Epoch 6	253	303	12
Epoch 4–Epoch 5	106	103	13
Epoch 4–Epoch 6	229	277	14
Epoch 5–Epoch 6	123	174	15

**Note.** By taking the difference between all possible combinations of observation epochs, we probe source variability on 15 timescales, from  $\sim 1$  month to  $\sim 5.5$  years. Throughout this paper, we refer to these epochal differences as intervals. The “label” column refers to the electronic table of the entire SAGE-Var catalog, available from the NASA/IPAC Infrared Science Archive (IRSA; columns 32–61, Tables 3 and 4).

criteria. The remaining 67 sources typically show marginally variable behavior, with variability indices very close to, but not quite exceeding, our  $3\sigma$  level. We manually add to our catalog 66 of these 67 sources, omitting one source only detected in one of the four new epochs of SAGE-Var.

These criteria resulted in 2198 unique variables in the LMC along with 571 in the SMC. Histograms showing the number of sources seen as variable in each interval are shown in Figure 4. Intervals 1–9 show a generally higher number of detected variables because they all compare at least one of the initial two SAGE epochs to later observations. Taken during the cold *Spitzer* mission, which had a S/N approximately twice that obtained during the warm mission when all other epochs in SAGE-Var were taken. Interval 10 shows an unusually small number of variables because it is an order of magnitude shorter than any other interval, spanning only  $\sim 20$  days. A randomly phased interval that short is unlikely to catch variability for many IR sources.

In order to classify the variables we did detect, we matched our detected variables against the OGLE-III Catalog of Variable Stars (CVS),<sup>13</sup> which consists of  $\sim 150,000$  classified variables in the LMC ( $\sim 26,500$  in the SMC) with well characterized variability information. Again using a  $2''$  matching radius, we find 1361 OGLE-CVS matches to our 2198 SAGE-Var variables in the LMC, and 323 matches to our 571 variables in the SMC. These matches are shown in the [3.6]–[4.5] versus [4.5] CMD in Figure 5. The OGLE populations we find in our data are detailed in Tables 6 and 7, along with the appropriate reference to the relevant OGLE-CVS document, if any. See the cited references for thorough discussion and definitions of the various classifications. The OGLE observations were taken primarily in the *I* band, with some additional images in *V* and *B*. The OGLE camera has a pixel scale of  $0''.417/\text{pixel}$ , and concentrates 80% of the light in a

$0''.5$  disk. See Udalski et al. (1997) for a thorough discussion of the OGLE hardware and observing protocols.

LPVs, mainly evolved AGB stars, are the most numerous OGLE sources we detect in our sample. This is not surprising, as AGB stars are among the brightest objects in the IR sky, and many of our intervals probe timescales on which LPVs are expected to vary. Sources identified as LPVs previously have been matched to variability data from the MACHO survey (Riebel et al. 2010). The MACHO survey utilized two non-standard bands; a red ( $\sim 690 \mu\text{m}$ ) and a blue ( $\sim 520 \mu\text{m}$ ). The MACHO optics have a pixel scale of  $0''.635/\text{pixel}$  and a median PSF of  $2''$ . Details of the MACHO project and its photometry can be found in Alcock et al. (1997) and Alcock et al. (1999).

We detect OGLE variables of every other OGLE classification besides LPVs as well, down to nearly the limiting magnitude of our survey. These other classes tend to be blue in the diagrams of Figure 5, as the instability strip they occupy on the HR diagram places the peak of their SEDs blueward of the *Spitzer* bands.

We also matched our data to the list of Young Stellar Object candidates (YSOs) in the LMC (SMC) compiled by Carlson et al. (2012; Sewilo et al. 2013). We find 500 (337) YSOs in the entire SAGE-Var LMC (SMC) data set, but only 12 (4) of them are identified as variables using our criteria. These are also shown in Figure 5.

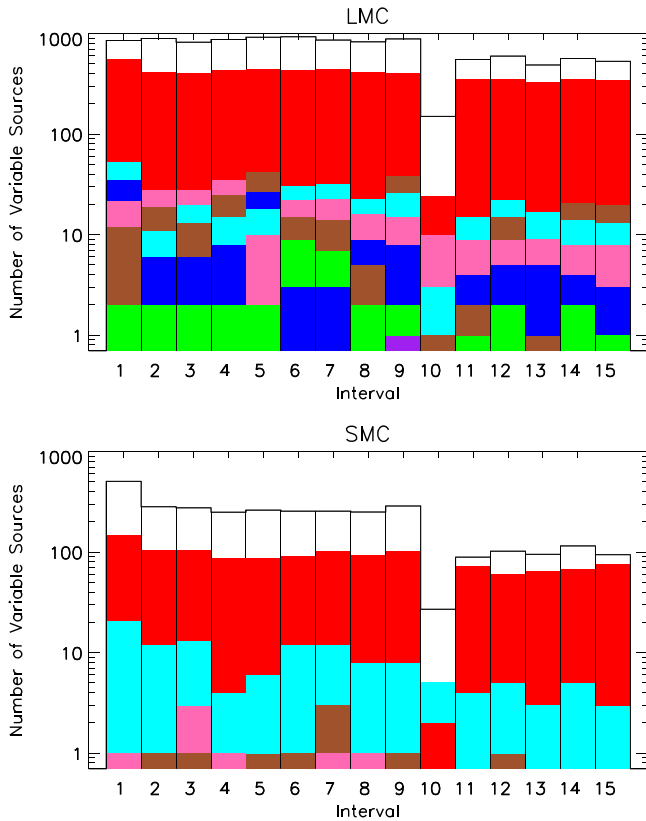
Figure 6 shows the same CMDs as Figure 5 but highlighting the sources without counterparts in the OGLE-CVS or MACHO survey. We remove any source with a SAGE [4.5] mosaic photometry dimmer than 15th magnitude identified as a variable but without an OGLE or MACHO counterpart. We include these sources in the online catalog for completeness, but data artifacts (such as blending with nearby sources) cause problems with our variability criteria for such dim sources. After removing those dim sources, we are left with 641 (139) IR variables in the LMC (SMC) without OGLE or MACHO detected variation. There are a few extremely red sources in the AGB region of the CMD without OGLE identifications. Most of these sources are flagged as probable variables by the *WISE* survey, and our results bolster that conclusion (Section 3.1).

### 3. RESULTS

#### 3.1. New AGBs

Because they typically require large amounts of telescope time and the expense of space-based multi-epoch observations often makes such lengthy surveys cost-prohibitive, most variability surveys are conducted from the ground, and hence at optical wavelengths. Ground-based variability surveys thus miss the reddest, most extreme AGBs, exactly the stars whose variability is most relevant to the evolved star dust budget because they dominate the mass return from AGBs to the ISM (Matsuura et al. 2009; Boyer et al. 2012; Riebel et al. 2012). SAGE-Var represents the first large-scale variability survey at such red wavelengths. As such, we have the ability to detect variable stars which have never been categorized as such before. Our survey detects variability in 641 (139) sources in the LMC (SMC) that do not have well-defined variability measurements in the OGLE or MACHO surveys (Figure 6). Out of these newly identified variables, 10 sources in the LMC and 6 in the SMC can be classified as AGB candidates, based on previous studies (Riebel et al. 2012), or their position in the *J*–*K<sub>s</sub>* versus *K<sub>s</sub>* CMD. Figure 7 highlights these 10 (6) new

<sup>13</sup> <http://ogledb.astrouw.edu.pl/~ogle/CVS/>



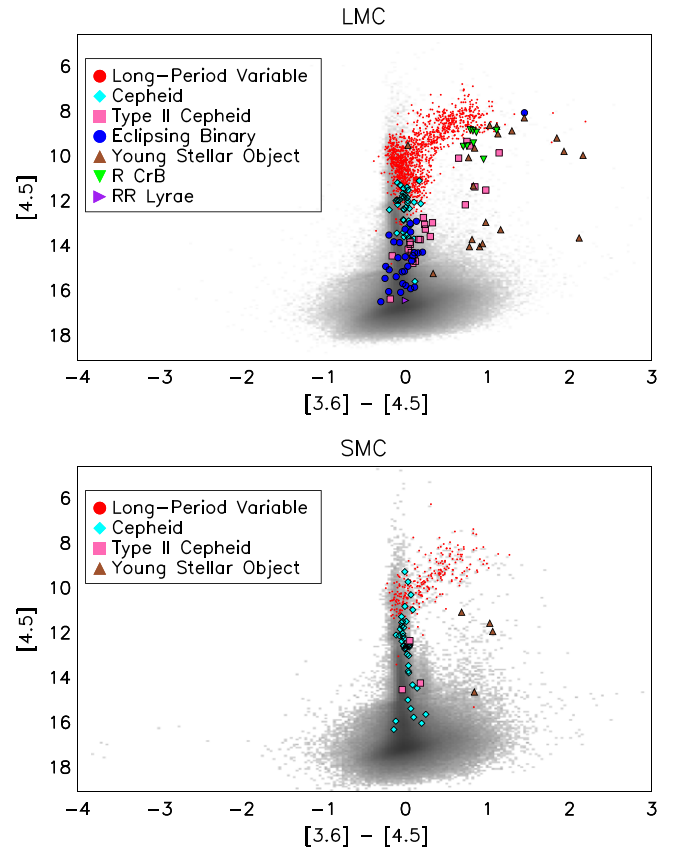
**Figure 4.** Histogram of the number of sources detected as variable in each interval of SAGE-Var. The top panel shows the source distribution in the LMC, while the bottom panel is that for the SMC. The colored portion of each bar represents the number of sources classified by the OGLE project. Red: LPVs, Cyan: Classical Cepheids, Hot Pink: Type II Cepheids, Blue: Eclipsing Binaries, Brown: YSOs, Green: R CrB stars, Purple: RR Lyrae stars. The white area at the top of each interval represents the unclassified variables seen in SAGE-Var but not listed in the OGLE-CVS (Figure 6). The labels for the intervals on the x-axis refer to the indices in Table 5. Classes are plotted from least numerous to most numerous, bottom to top.

AGB candidates against a Hess diagram of the entire SAGE-Var sample.

### 3.1.1. LMC AGB Candidates

We find 10 AGB candidates in the LMC without previously well-measured periods in either the OGLE or MACHO surveys. Most of these were flagged as highly likely to be true variables by the *WISE* survey, and we confirm that measurement. We do not find any of these 10 candidates in the SAGE-Spec (Kemper et al. 2010) list of LMC sources. Without a chemical classification derived from observed spectra, we therefore employ an Spectral Energy Density (SED)-based chemical classification.

We fit the SEDs of these sources with synthetic SEDs from the GRAMS model grid (Sargent et al. 2011; Srinivasan et al. 2011). GRAMS is a grid of dusty AGB/RSG star models computed by the radiative transfer code 2Dust (Ueta & Meixner 2003). The grid is constructed by placing spherically symmetric circumstellar dust shells of varying inner radii, optical depth, and both O-rich and C-rich chemical composition around model stellar photospheres. The models assume a constant mass-loss rate (i.e., an inverse-square density law) in the shell and a power-law distribution of grain sizes with an exponential falloff, with a typical size of  $\sim 0.1 \mu\text{m}$ . The O-rich



**Figure 5.**  $[3.6] - [4.5]$  vs.  $[4.5]$  CMD highlighting the OGLE-CVS variables with counterparts among the SAGE-Var classified variables. The LMC sources are shown in the top panel and the SMC sources are shown in the bottom. The entire SAGE-Var catalog is shown as a grayscale Hess diagram in the background. The same color scheme as in Figure 4 is used.

dust is modeled using the astronomical silicates of Ossenkopf et al. (1992), and the carbonaceous dust is a 9:1 mixture by mass of amorphous carbon (optical constants from Zubko et al. 1996) and silicon carbide (optical constants from Pegourie 1988). We perform a minimum  $\chi^2$  fitting to obtain the best-fit parameters, which include the luminosity, dust-production rate, and chemical type.

Using the GRAMS model grid, we classify 8/10 of these AGB candidates as C-rich. C-rich AGB stars tend to be redder than O-rich stars, and our newly measured variables skew red. AGB stars in our sample detected by the OGLE survey have an average  $[3.6] - [4.5]$  color of  $-0.07$  mag, while the sources with no OGLE detection are nearly a magnitude redder at these wavelengths, with an average  $[3.6] - [4.5]$  of  $0.74$  mag. SAGE-Var has too sparse a sampling of the light curves to deduce a period for their variability, but we can place lower limits on their IR variability amplitudes. These sources are listed in Table 8.

### 3.1.2. SMC AGB Candidates

After cross-matching our list of AGB candidates in the SMC with SIMBAD and removing all sources confirmed to not be AGB stars, we generate a list of six new AGB candidates. We verified that none of these candidates were in the P. Ruffe et al. (2015, in preparation) list of spectroscopically classified SMC sources. Using the GRAMS model grid, we classify five out of six of them as C-rich. These sources are listed in Table 9.



**Table 6**  
Matched Variable Populations Detected in SAGE-Var LMC

Variable Classification	Number	Reference/Definition
Total Variables	2198	
Detected		
Long-period Variables	1065	Soszyński et al. (2009b)
AGB C	5	...
AGB O	6	...
Mira AGB C	426	...
Mira AGB O	143	...
OSARG AGB C	15	OGLE small-amplitude variable reg giant (OSARG)
OSARG AGB O	42	...
OSARG RGB C	1	...
OSARG RGB O	17	...
SRV AGB C	262	Semi-regular variable (SRV)
SRV AGB O	148	...
Cepheids	28	Soszyński et al. (2008a)
1O	1	1st overtone pulsation mode
1O/2O	1	Mixture of 1st and 2nd overtone pul- sation modes
F	26	Fundamental-mode pulsation
Type II Cepheids	19	Soszyński et al. (2008b)
BLHer	1	...
RVTau	12	...
WVir	6	...
RR Lyrae	3	Soszyński et al. (2009a)
R CrB	6	Soszyński et al. (2009c)
Eclipsing Binaries	25	Graczyk et al. (2011)
EC	8	Contact binaries
ECL	1	Ellipsoidal contact binaries
ED	8	Detached binaries
ESD	8	Semi-detached binaries
Young Stellar Objects	12	Carlson et al. (2012)

**Note.** This table lists the variable classifications and subclasses used by the OGLE Catalog of Variable Stars and the YSO catalog of Carlson et al. (2012), and the numbers of such sources also detected as variable by the SAGE-Var survey in the LMC. A  $2''$  matching radius was used for all OGLE catalog comparisons, and the YSO list was matched based on SAGE Archive designations. See the cited references for complete definitions of the subcategories listed here. Uncommon abbreviations are defined in the third column.

Source SSTISAGEMA J010041.61–723800.7 is the reddest new LPV candidate we identify in the SMC. The best-fitting GRAMS model is an O-rich model, but the extreme redness of the source is more consistent with a C-rich star. Currently, the GRAMS model grid in the SMC places too much weight on the  $24\ \mu\text{m}$  photometry, and this leads to some sources being misclassified (see S. Srinivasan et al. 2015, in preparation for details). We manually change the classification of this source to C-rich.

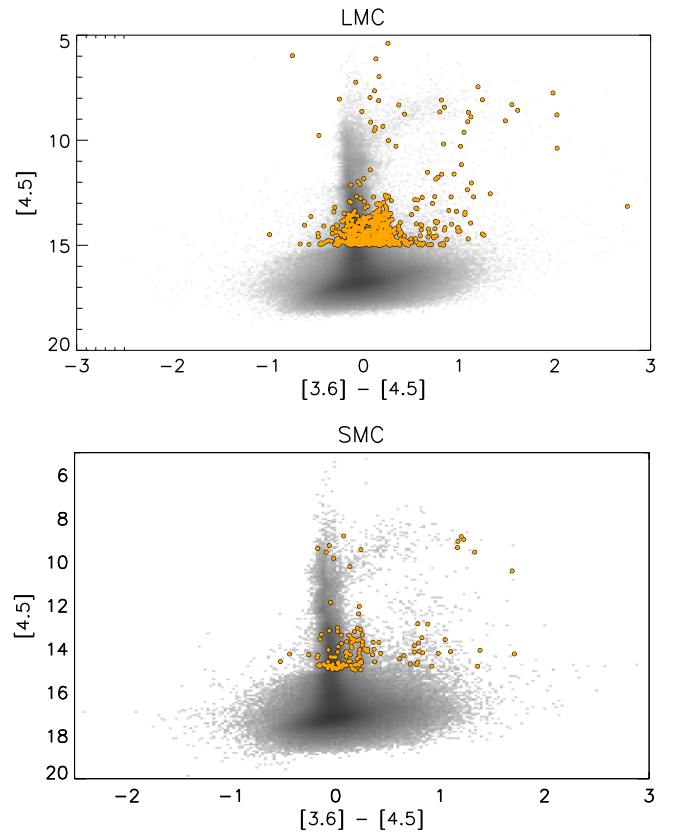
### 3.2. Variability Amplitude and Dust Production Rate

Le Bertre (1992) found a strong correlation between variability amplitude (measured in the  $K$  band), and  $K - L'$  color for a sample of 20 carbon-rich LPVs in our galaxy. Whitelock et al. (1991) and Le Bertre (1993) found a similar correlation for O-rich LPVs. Whitelock et al. (1994) found a strong correlation between the IR amplitude and the  $K - [12]$  color. The  $K$  band is dominated by light from the stellar photosphere, while the  $IRAS\ 12\ \mu\text{m}$  band represents the emission from the cooler dust shell around the star, and so this color

**Table 7**  
Matched Variable Populations Detected in SAGE-Var SMC

Variable Classification	Number	Reference
Total Variables Detected	571	...
Long-period Variables	276	Soszyński et al. (2011)
Mira C	140	...
Mira O	22	...
OSARG C	7	...
OSARG O	4	...
SRV C	95	...
SRV O	8	...
Cepheids	42	Soszyński et al. (2010a)
F	40	...
1O	2	...
Type II Cepheids	3	Soszyński et al. (2010b)
RVTau	1	...
WVir	1	...
pWVir	1	...
Young Stellar Objects	4	Sewiło et al. (2013)

**Note.** This table lists the variable classifications and subclassifications used by the OGLE Catalog of Variable Stars and the YSO catalog of Sewiło et al. (2013), and the numbers of such sources also detected as variable by the SAGE-Var survey in the SMC. A  $2''$  matching radius was used for all OGLE catalog comparisons, and the YSO list was matched based on SAGE Archive designation.



**Figure 6.**  $[3.6] - [4.5]$  vs.  $[4.5]$  CMD highlighting variables detected by SAGE-Var but not identified by the OGLE or MACHO surveys. The LMC sources are shown in the top panel, and the SMC sources are shown in the bottom. The entire SAGE-Var catalog is shown as a grayscale Hess diagram in the background. To eliminate false positives due to low S/N, we only consider sources brighter than 15th magnitude to be variables.

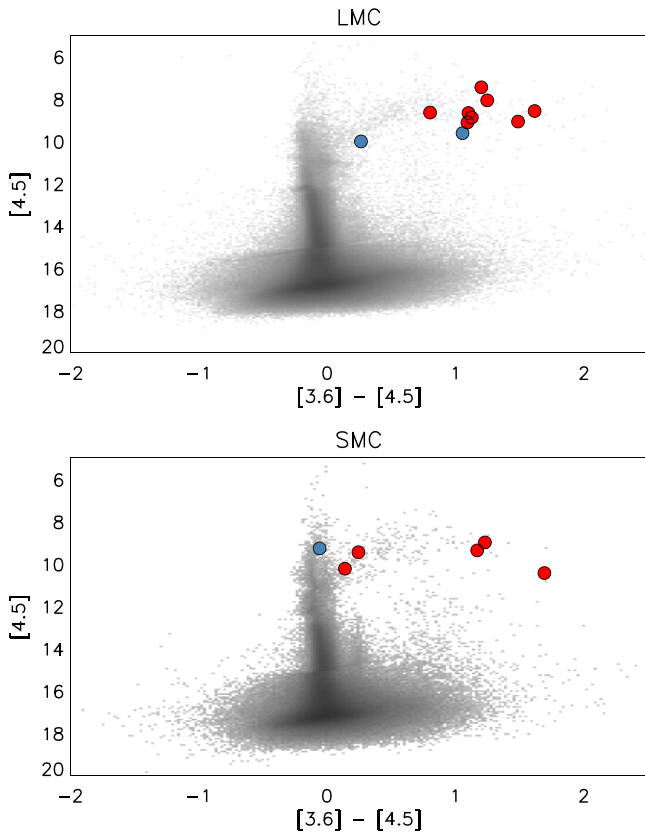
**Table 8**  
AGB Candidates in the LMC without OGLE or MACHO Variability Measurement

SAGE Designation <sup>a</sup>	R.A. (2000)	Decl. (2000)	GRAMS Class	[3.6] Amplitude	[4.5] Amplitude	[3.6]–[4.5] Color
J051041.21-683606.6	77.6717	-68.6018	C <sup>b</sup>	0.83	0.85	1.24
J051414.85-700409.8	78.5619	-70.0694	C <sup>b</sup>	0.93	0.86	1.12
J052503.26-692617.3	81.2636	-69.4381	C <sup>b</sup>	0.90	0.69	1.48
J052813.02-691228.4	82.0543	-69.2079	C	0.86	0.75	1.09
J052900.19-695247.3	82.2508	-69.8798	C	1.11	0.52	0.80
J053051.75-694328.0	82.7156	-69.7245	C	1.42	1.19	1.20
J050202.38-690726.2	75.5099	-69.1239	C	0.52	0.56	1.10
J050718.89-683850.4	76.8287	-68.6474	C	1.50	1.27	1.61
J051913.89-693818.3	79.8079	-69.6384	O	0.29	0.31	0.26
J053010.30-690933.8	82.5429	-69.1594	O	0.36	0.30	1.05

**Notes.** The 10 AGB candidates in the LMC with variability newly detected by SAGE-Var. Many of these sources are confirmed to be AGB stars in the literature, but none have previously been observed to vary. The GRAMS Class column lists the classification (O-rich or C-rich) assigned each source by the GRAMS model grid. The next two columns represent a lower bound on the IR variability amplitude of the sources, the maximum observed magnitude minus the minimum observed magnitude. The [3.6]–[4.5] color is also listed to connect the entries in this table to Figure 7.

<sup>a</sup> Designations in the online data table are prefaced with “SSTISAGEMA”

<sup>b</sup> These three sources have been spectroscopically confirmed as C-rich stars by van Loon (2006).



**Figure 7.** [3.6]–[4.5] vs. [4.5] CMD highlighting newly identified LPV candidates. The LMC sources are shown in the top panel, and the SMC sources are shown in the bottom. The entire SAGE-Var catalog is shown as a grayscale Hess diagram in the background. Stars classified as C-rich by the GRAMS model grid are shown in red, stars classified as O-rich are shown in blue.

serves as an indicator of the thickness of the circumstellar dust shell, and hence the rate of dust production by the star, if DPR is assumed to be constant.

By using the GRAMS modeling results of Riebel et al. (2012), we can directly compare the [3.6] amplitude and the DPR for our sample. Our results are shown in Figure 8. We

extract every source in the SAGE-Var catalog with a valid GRAMS classification which was also classified as a true variable (Section 2.4). We see a slight correlation of increasing DPR with increasing IR amplitude (albeit one with a very large scatter). C-rich and O-rich AGB stars have considerable overlap in their range of DPR, but more C-rich stars than O-rich stars extend to higher DPRs (see discussion in Riebel et al. 2012). Because of this, the trend is more visible among the C-rich sources, but both populations follow the same basic trend, with considerably more scatter at low amplitudes for the O-rich sources.

### 3.3. PL Relationships

#### 3.3.1. AGB PL Relationships

Characterizing the AGB PL relationship is important in light of the upcoming launch of the *James Webb Space Telescope* (JWST). AGB stars are among the brightest sources in the infrared (IR) sky, and with a large orbital platform concentrating on the IR, AGB stars could serve as important distance indicators, if their intrinsic luminosity can be well determined from their variability. Previous investigations of the AGB PL relationship (see Riebel et al. 2010, and references therein) have found a scatter of a few tenths of a magnitude about the best-fit line, and it has not been clear whether this scatter was intrinsic to the relationship (i.e., a real phenomenon due to an astrophysical cause such as a spread in stellar masses) or an artifact resulting from observing stars at various points in their light curves. Constructing mean magnitudes from observations taken at several points in an AGB star’s light curve are an essential step toward clarifying the nature of this spread. The 6 epochs of the SAGE-Var survey allow us to calculate improved mean magnitudes for every source in the catalog.

We use only AGB stars from the sample of Riebel et al. (2010) that are identified as variables in the SAGE-Var data to construct our relations. Only stars that have been observed at more than one brightness benefit from the averaging process, as opposed to stars with six measurements of the same brightness. For all AGB stars identified as variables, all available valid flux measurements from the SAGE and SAGE-Var surveys were



**Table 9**  
AGB Candidates in the SMC without OGLE or MACHO Variability Measurement

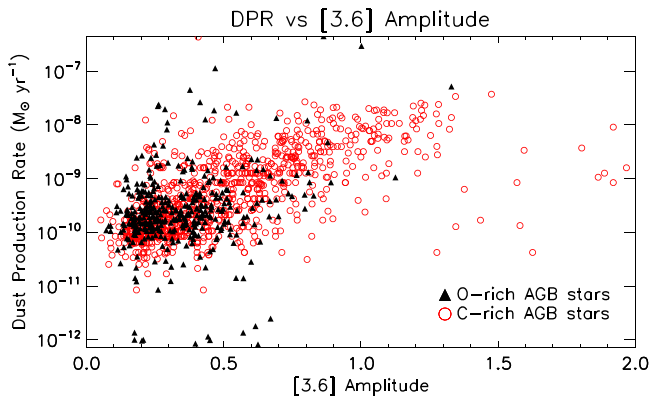
SAGE Designation <sup>a</sup>	R.A. (2000)	Decl. (2000)	GRAMS Class	[3.6] Amplitude	[4.5] Amplitude	[3.6]–[4.5] Color
J005106.28-731635.9 <sup>b</sup>	12.7762	-73.2767	C	0.34	0.34	0.14
J004544.12-720815.4	11.4338	-72.1376	C	0.32	0.40	0.24
J005926.35-722341.4	14.8598	-72.3949	C	1.09	0.88	1.23
J010232.75-721912.5	15.6365	-72.3202	C	1.08	1.06	1.17
J010041.61-723800.7	15.1734	-72.6335	C <sup>c</sup>	0.95	1.03	1.69
J005131.21-732007.7	12.8801	-73.3355	O	0.34	0.23	-0.06

**Notes.** The 6 AGB candidates in the SMC with variability newly detected by SAGE-Var. Most of these sources are confirmed AGB stars, but none have previously been observed to vary. The GRAMS Class column lists the classification (O-rich or C-rich) assigned each source by the GRAMS model grid (S. Srinivasan, et al., 2015, in preparation). The next two columns represent a lower bound on the IR variability amplitude of the sources, the maximum observed magnitude minus the minimum observed amplitude. The [3.6]–[4.5] color is also listed to connect the entries in this table to Figure 7.

<sup>a</sup> Designations in the online data table are prefaced with “SSTISAGEMA.”

<sup>b</sup> The SAGE coordinates for this source are approximately 2″ away from the SIMBAD source ISO-MCMS J005106.5-731636, discussed by Cioni (2003), but visual inspection of the images indicates they may be the same object

<sup>c</sup> The best-fitting GRAMS model for this source is an O-rich model with poor fit quality. Based on the extreme redness of this source, we classify it as C-rich



**Figure 8.** Dust production rate (DPR,  $M_{\odot} \text{ yr}^{-1}$ ) vs. [3.6] amplitude for AGB stars in the LMC and SMC. The DPR is taken from the best-fitting GRAMS model, and the [3.6] amplitude is defined as the dimmest measured SAGE-Var [3.6] magnitude minus the brightest [3.6] magnitude. C-rich AGB stars are shown as open red circles, and O-rich AGB stars as filled black triangles.

averaged to construct mean fluxes at both [3.6] and [4.5]. Photometric errors from each observation were added in quadrature to produce an error for the mean flux. These mean fluxes were transformed into magnitudes using the zero points given in Section 2.2.

Figure 9 shows the PL relations constructed from the six-epoch mean magnitudes constructed from the SAGE-Var data. The 3.6  $\mu\text{m}$  magnitude is used in the top panel, and the 4.5  $\mu\text{m}$  band is used in the bottom. The periods come from the MACHO survey if possible, and from the OGLE survey if no MACHO period was available. The MACHO periods were given priority in order to maintain consistency with the study of Riebel et al. (2010). For those stars with measured periods in both the OGLE and MACHO surveys, the periods were found to differ by only three days on average. We only present updated fits for stars pulsating in the fundamental mode (sequence 1), as this was the only sequence with a significant population observed in SAGE-Var. The derived PL relations for the stars classified as O-rich (by the best-fitting GRAMS model) are shown in Figure 9 and quantitatively described in

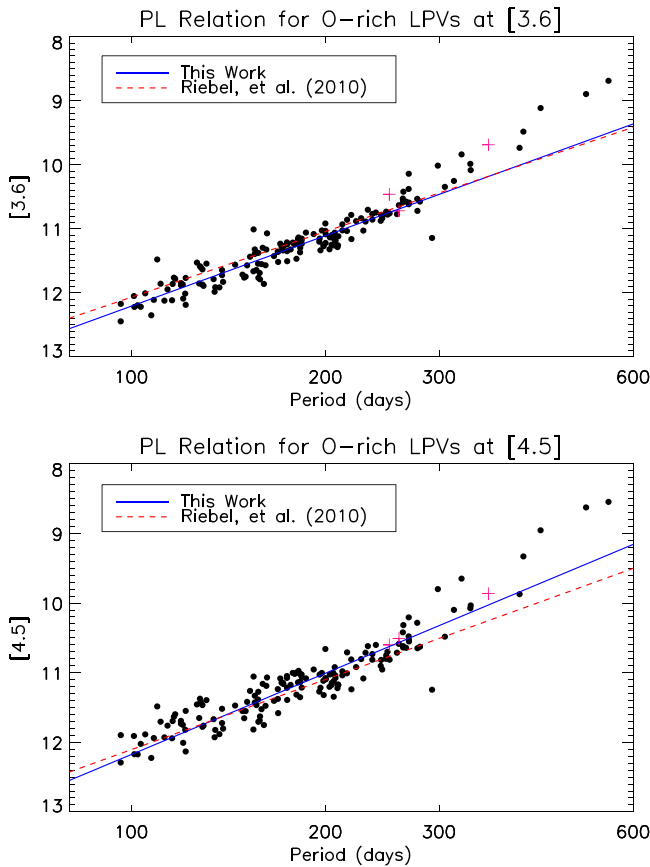
Table 10. The fits to the stars identified as C-rich are shown in Figure 10 and described in Table 11.

In Figure 9 stars with periods  $\gtrsim 300$  days are systematically brighter than the trend line, whereas in Figure 10 stars with periods  $\lesssim 300$  days are systematically dimmer than the trend line. These outlier stars are an overlap between the two populations. The C-rich best-fit trend line is about a magnitude brighter than the O-rich fit, and C-rich stars have longer periods on average. Thus, the shortest-period C-rich stars that are dimmer than the trend line actually overlap the main body of the O-rich stars in PL space. In the same manner, the longest period O-rich stars which are brighter than the trend line in Figure 9 overlap the main population of C-rich stars in Figure 10. When using photometric classification, the division between O- and C-rich AGB stars is not sharp, and the two populations smoothly blend into one another in PL space (see Figures 1 and 5 in Riebel et al. 2010). Because there are only about 10 of these outlying stars, the fits are insensitive to their inclusion, and they have not been removed from the sample prior to fitting.

Using the criteria described in Appendix A of Riebel et al. (2010), we identify three O-rich AGB stars in the SMC which lie on Sequence 1. We corrected for the difference in distance between the LMC and the SMC using distance moduli of 18.54 and 18.93, respectively (Keller & Wood 2006). These stars, shifted to the distance of the LMC, are plotted as pink crosses in Figure 9. On average, the LMC has a greater metallicity than that of the SMC. Our numbers are too small to draw definite conclusions, but we do not see any evidence for a dependence of the AGB PL relation on metallicity in our sample. This idea is worthy of further investigation.

While we do see a slight decrease in the scatter about the best-fit line compared to that found by Riebel et al. (2010), the reduction is less than a factor of two, which leads us to believe that much of this remaining scatter is intrinsic to the relationship and is not a product of observing different stars at different phases of their light curve.

We identified 223 (49) Carbon-rich, non-extreme AGB stars pulsating in the fundamental mode in the LMC (SMC). “Extreme” AGB stars are defined as those stars with  $J - [3.6] > 3.1$ . The flux in the  $J$  band traces the emission



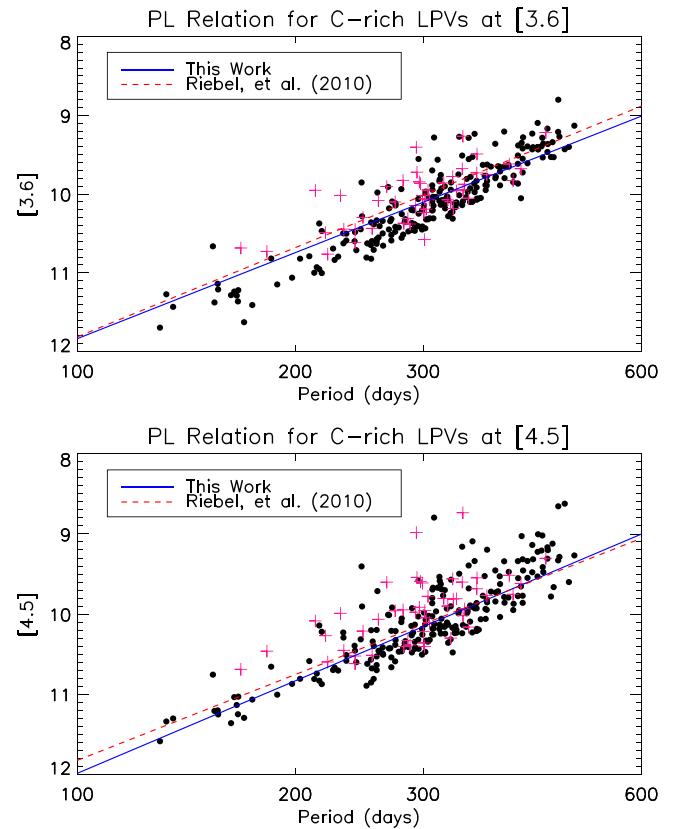
**Figure 9.** PL relationships for GRAMS classified O-rich AGB stars in the SAGE-Var sample, constructed using six-epoch mean magnitudes in the [3.6] and [4.5] bands. Stars from the LMC are shown as black points; stars from the SMC shown as pink plus signs. Note that the horizontal axis is a log scale.

**Table 10**  
Period–Magnitude Relationships for O-rich AGB Stars  
Pulsating in the Fundamental Mode

PL Relation	$N$	Scatter (mag)
This work:		
[3.6] = $-4.00(\pm 0.03) \log P + 20.24(\pm 0.07)$	131	0.18
[4.5] = $-3.78(\pm 0.03) \log P + 19.72(\pm 0.06)$	131	0.22
Riebel et al. (2010):		
[3.6] = $-3.41(\pm 0.04) \log P + 18.88(\pm 0.09)$	2221	0.271
[4.5] = $-3.35(\pm 0.04) \log P + 18.80(\pm 0.01)$	2227	0.270

**Note.** PL relations for LPVs classified as O-rich AGB stars. The quoted scatter is the standard deviation of the residuals about the best-fit line.

from the stellar photosphere, while the [3.6] emission is from the circumstellar dust shell around these heavily enshrouded stars. At SAGE-Var wavelengths, these stars do not follow the same PL relation as less enshrouded stars because their brightness does not reflect stellar emission, but emission from the dust shell (see Figure 4 in Riebel et al. 2010). Sources thus classified as extreme are not included in this fit. The derived PL relations to the remaining C-rich stars are shown in Figure 10 and quantitatively described in Table 11. The [4.5] data are shown for completeness, but we do not attribute any significance to the linear fits in that band. As discussed by Riebel et al. (2010), AGB stars are affected by a CO absorption



**Figure 10.** PL relationships for GRAMS classified C-rich AGB stars in the SAGE-Var sample, constructed using six-epoch mean magnitudes in the [3.6] band. Stars from the LMC are shown as black points; stars from the SMC shown as pink plus signs. Note the horizontal axis is a log scale.

feature in the [4.5] band, which distorts the linear PL relationship in that band. O-rich stars do not have enough CO in their atmospheres to be greatly affected. Figure 6 in Riebel et al. (2010) shows that the slope of the O-rich AGB PL relation is relatively insensitive to wavelength, whereas the C-rich relation shows significant variation.

Correcting for distance, the SMC stars are well-mixed with the LMC stars, showing comparable scatter about the best-fit line. With many more samples than the O-rich, this provides more substantial evidence that C-rich AGB stars in the SMC follow the same PL relationship as those in the LMC.

### 3.3.2. Cepheid IR PL Relationships

The well-known PL relationship of classical Cepheid variables (the Leavitt Law) is one of the most important extragalactic distance indicators in astronomy. Cepheid studies have generally focused on the optical, but McGonegal et al. (1982) pointed out that the near-IR offers several advantages for calibration of the Cepheid PL relation. Interstellar reddening, which imposes an intrinsic scatter on magnitude determinations in the optical, is much smaller at *Spitzer* wavelengths. Also, Cepheid variability amplitudes are considerably smaller in the IR, and thus sparsely sampled lightcurves are good tracers of the mean brightness. For these reasons, several recent surveys have focused on calibrating the Cepheid PL relation in the IR (Scowcroft et al. 2011; Freedman et al. 2012; Ngeow et al. 2012).

**Table 11**  
Period–Luminosity Relationships for C-rich AGB Stars Pulsating in the Fundamental Mode

PL Relation	$N$	Scatter (mag)
This work:		
[3.6] = $-3.63(\pm 0.02) \log P + 19.11(\pm 0.05)$	272	0.24
[4.5] = $-3.84(\pm 0.02) \log P + 19.66(\pm 0.06)$	272	0.30 Riebel et al. (2010):
[3.6] = $-3.77(\pm 0.05) \log P + 19.35(\pm 0.12)$	1813	0.251
[4.5] = $-3.56(\pm 0.05) \log P + 18.96(\pm 0.12)$	1816	0.265

**Note.** PL relations for LPVs classified as C-rich AGB stars. The quoted scatter is the standard deviation of the residuals about the best-fit line.

**Table 12**  
Classical Cepheid Period–Luminosity Relationships at  
[3.6] and [4.5] in the Magellanic Clouds

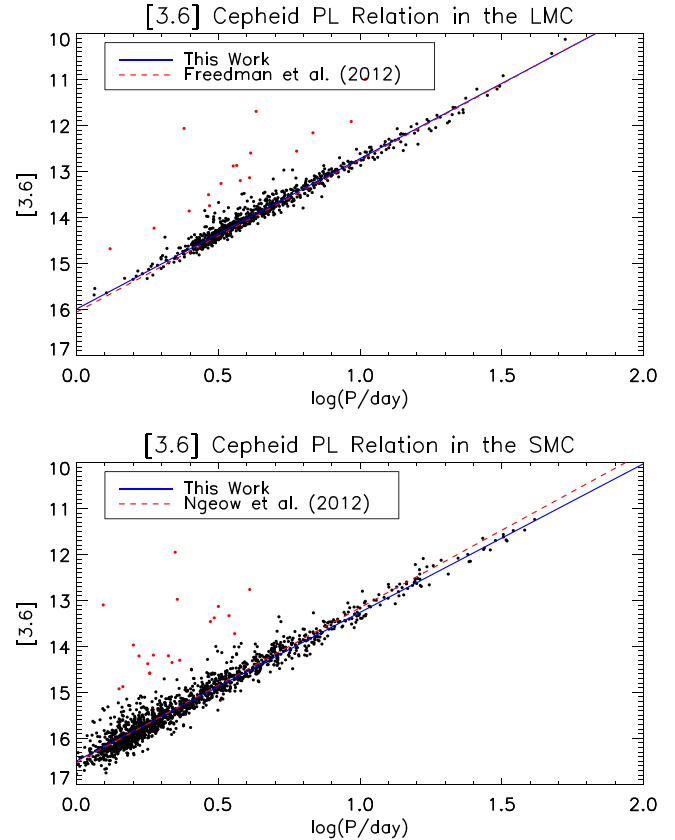
PL Relation	$N$	Scatter (mag)
LMC:		
[3.6] = $-3.271(\pm 0.004) \log P + 15.993(\pm 0.003)$	811	0.13
[4.5] = $-3.157(\pm 0.004) \log P + 15.877(\pm 0.002)$	820	0.13
SMC:		
[3.6] = $-3.261(\pm 0.006) \log P + 16.511(\pm 0.004)$	452	0.18
[4.5] = $-3.437(\pm 0.004) \log P + 16.665(\pm 0.002)$	454	0.18

**Note.** PL relations for fundamental mode classical Cepheids in the LMC and SMC. Fits were determined using a standard  $3\sigma$  clipping procedure. The fits to the SMC data were determined by only considering Cepheids with  $\log P > 0.5$ . The scatter about the fit is defined to be the standard deviation of the residuals.

These studies have focused on small ( $\sim 80$ ) samples of Cepheids, with well sampled light curves. SAGE-Var was designed as a more general variability survey, so what we lack in thorough coverage of the light curve, we make up for in increased sample size. Based on OGLE classification, we find 837 (1536) fundamental mode Cepheids in our LMC (SMC) data.

We extract every source observed by SAGE-Var and classified as a fundamental mode Cepheid by the OGLE survey (Cepheids primarily varying in the first or higher overtone mode will by definition not lie on the primary PL relation). We eliminate any source with only one valid observation during SAGE-Var, as this provides limited means to constrain the uncertainty on the mean flux of the source. We calculate a simple mean magnitude for all our sources by averaging the fluxes of all SAGE-Var observations and then converting to magnitudes. Only four of our sources are in common with Scowcroft et al. (2011), and our mean magnitudes derived from randomly phased observations are within 0.05 magnitudes of theirs. We use the periods from the OGLE survey (Soszyński et al. 2008a, 2010a).

The least-squares linear fits (using  $3\sigma$  clipping) are given below in Table 12. The Leavitt Law at [3.6] derived from the LMC sample is illustrated in Figure 11 and the relation at [4.5] is shown in Figure 12. Fits are of the form  $y = A \log P + B$ , with  $y$  the mean magnitude of the source, and the period  $P$  measured in days. The scatter is defined as the standard deviation of the residuals to the fit. We visually inspected all the sources clipped as part of the fitting process. More than 90% are obvious blending/confusion issues in the original data. We find that when fitting the SMC data, using only sources with  $\log P > 0.5$  provided a better visual fit to the data (this method was also used by Scowcroft et al. 2011). The LMC fits

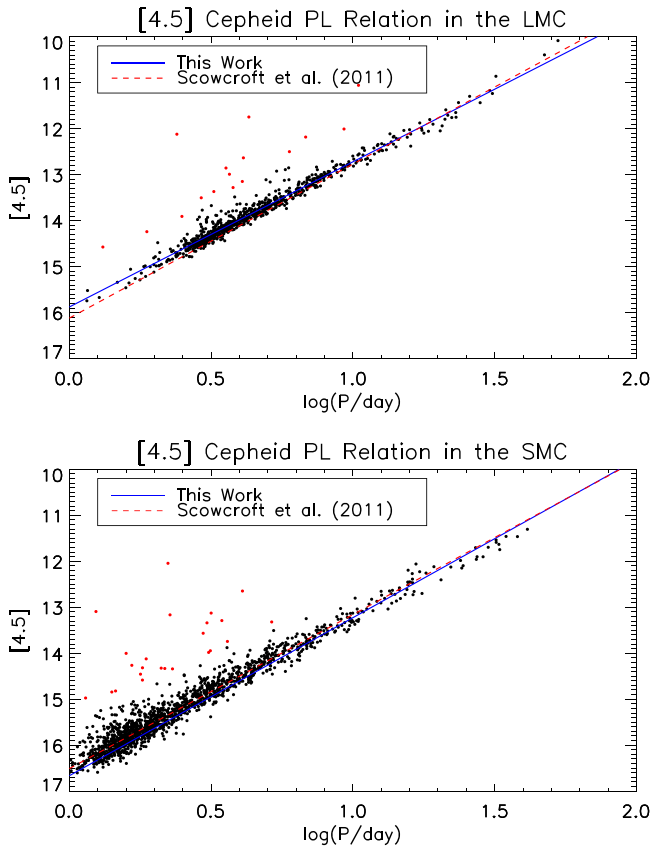


**Figure 11.** 3.6  $\mu\text{m}$  Leavitt Law for 811 (452) classical Cepheids in the LMC (SMC). The LMC is shown in the top panel and the SMC in the bottom. The fits are quantitatively detailed in Table 12. Stars shown in red have residuals to the fit greater than  $3\sigma$  and did not contribute to the fit. The relation derived in this study is plotted in blue. For the LMC (SMC) data, the relation determined by Freedman et al. (2012; (Ngeow et al. 2012) is overplotted in red as a dashed line.

were robust to this decision, and no period selection criteria were applied. The results given in Table 12 show the greater scatter about the SMC relationship compared with that in the LMC. This is consistent with the much greater overall line-of-sight depth of the SMC compared with LMC seen in the early Cepheid work of Caldwell & Coulson (1986), as well as in other stellar components. This is probably due to gravitational interactions in the Magellanic system (see Scowcroft et al. 2015, and references therein).

The slope of our [3.6] relations show very good agreement with those found by Freedman et al. (2012) in the LMC and by Ngeow et al. (2012) in the SMC. In the [4.5] band, we overplot the relations determined by Scowcroft et al. (2011), which also show good agreement with our own. As discussed by Scowcroft et al. (2011) and Monson et al. (2012), CO absorption in the [4.5] band renders this PL relation





**Figure 12.**  $4.5\ \mu\text{m}$  Leavitt Law for 811 (452) classical Cepheids in the LMC (SMC). The LMC is shown in the top panel and the SMC in the bottom. The fits are quantitatively detailed in Table 12. Stars shown in red have residuals to the fit greater than  $3\sigma$  and did not contribute to the fit. The relation derived in this study is plotted in blue. The relation determined by Scowcroft et al. (2011) using the original SAGE-LMC data is overplotted in red as a dashed line.

problematic for Cepheid distance determinations. This effect is clearly seen in Table 12, where the slopes of the PL relations at  $[3.6]$  in the the LMC and SMC agree well, while the slopes at  $[4.5]$  do not. This seems likely to be due to the different effects of CO absorption at different metallicities.

#### 4. CONCLUSIONS

We present the results from a four-epoch unbiased IR survey of the central regions of the LMC and SMC. We have produced full catalogs of our observations, consisting of 1,717,554 (457,760) objects in the LMC (SMC). We have identified 2198 (571) objects in the LMC (SMC) as probable variables.

We identify 10 (6) variable AGB candidates in the LMC (SMC) without well-determined variable periods from OGLE or MACHO (Section 3.1). Most of these sources were flagged as potential variables in the *WISE* survey, and our independent measurement confirms that probability.

Using mean magnitudes constructed from six epochs of observations, we investigate the PL relationship for AGB stars pulsating in the fundamental mode. We find no significant reduction in the scatter about the best-fit relation compared to the results of Riebel et al. (2010), indicating the scatter ( $\sim 0.2$  mag) might be intrinsic to the data (Section 3.3.1). We find no evidence that the PL relationships of AGB stars in the SMC

( $Z \sim 0.04$ ) and in the LMC ( $Z \sim 0.08$ ) are different due to metallicity effects.

We present infrared PL relations for a sample of 811 (452) Cepheids in the LMC (SMC). Cepheid amplitudes are small in the IR, and our random-cadence results compare very well with those of Scowcroft et al. (2011), Freedman et al. (2012), and Ngeow et al. (2012; Section 3.3.2).

This publication made use of the VizieR database operated by CDS, Strasbourg, France. M.W.F. and P.A.W. gratefully acknowledge the receipt of research grants from the National Research Foundation (NRF) of South Africa. Meixner acknowledges support from NASA NAG5-12595, and *Spitzer* contract NM0710076. This work is based on observations made with the *Spitzer Space Telescope*, which is operated by the Jet Propulsion Laboratory, California Institute of Technology under a contract with NASA. M.L.B. is supported by the NASA Postdoctoral Program at the Goddard Space Flight Center, administered by ORAU through a contract with NASA. S.S. acknowledges support from the National Science Council and the Ministry of Science and Technology in the form of grant MOST103-2112-M-001-033-.

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